# **XMM-Newton Calibration Updates**

Michael Smith, ESAC IACHEC on-line meeting 23 May 2022

# XMM calibration releases over 2021/22

	XMM-CCF-REL-		
Astrometry:	time variable boresight updates	380, 387	
RGS:	bad pixels, gain, CTI	381, 383, 385	
EPIC:	effective area correction (EPIC-MOS - EPIC-pn)	382	
	effective area correction (EPIC-pn - NuSTAR)	388	(see talk by Felix Fuerst)
	EPIC-pn energy scale	389	(see talk by Ivan Valtchanov)
OM:	grism time dependent sensitivity	386	



### **RGS: contamination**



Contamination estimated from the flux of RX J1856-357 at long wavelengths.

Observed flux decrease not explained by  $C_8H_8$  contamination.

Empirical corrections in place.

Work ongoing to understand reason of Aeff loss.





#### **RGS: wavelength scale**

AB Dor, Capella, HR 1099: Revolution 2500 3000 3500 ABDOR CAPELLA 20 HR1099 10 0 -10 -20 Shift (mÅ) ABDOR CAPELLA 20 HR1099 10 0 -10 -20 2018 2020 2014 2016 2012 Year





#### **EPIC: effective area**

**MOS** Build up of contaminant, esp. MOS2

Based on 1E 0102.2-7219







## **EPIC: updated A<sub>eff</sub> correction**

Recalibration of the CORRAREA correction: an empirical correction of MOS A<sub>eff</sub> to PN

Sample of ~ 120 sources:		1.4
<ul> <li>On-axis, point source, non-piled up</li> </ul>		1.3
Per observation:	Z	1.2
Derive best-fit PN model	LL.	1.1
<ul> <li>Apply PN model to MOS1 &amp; MOS2</li> </ul>		1.0
Per instrument, stack		0.9 1.4 0.
<ul> <li>Data</li> <li>Model (= expected cts + scaled bkg)</li> </ul>		1.3
Determine stacked data / model ratios and normalise to PN	11	1.2
Determine stacked data / model ratios and normalise to PN	2	1.1
Derive energy-dependent A <sub>eff</sub> correction function (spline) to minimise residuals:		1.0
• correction to MOS $A_{eff} > 2.0 \text{ keV}$		0.9 1.4
<ul> <li>null correction &lt; 2.0 keV (where redistribution effects may</li> </ul>		1.3

be significant)



1.2

1.0

0.9

М2



# **EPIC: updated A<sub>eff</sub> correction**

Recalibration of the CORRAREA correction: an empirical correction of MOS A<sub>eff</sub> to PN

Sample of ~ 120 sources:

• On-axis, point source, non-piled up

Per observation:

- Derive best-fit PN model
- Apply PN model to MOS1 & MOS2

Per instrument, stack

- Data
- Model (= expected cts + scaled bkg)

Determine stacked data / model ratios and normalise to PN

Derive energy-dependent A<sub>eff</sub> correction function (spline) to minimise residuals:

- correction to MOS  $A_{eff} > 2.0 \text{ keV}$
- null correction < 2.0 keV (where redistribution effects may be significant)

M2

Σ



7

# **OM: grism time-dependent degradation**



Measurements of spectra of 3 standard stars in 6 wavelength bands in each grism.



Decline of sensitivity corrected in SAS; calibration updated to 2024 in (XMM-CCF-REL-386 released 03/12/21)

#### **OM grism throughput in 2024.0**

Grism	Throughput
UV	0.86
VIS	0.90



# **OM: monitoring time-dependent degradation (filters)**



#### **OM throughput at 2030**

Filter	Throughput
V	0.86
В	0.89
U	0.89
UVW1	0.86
UVM2	0.81
UVW2	0.78

Declines continue to slow in all filters.

Degradation is corrected in SAS.

